

V&V Summary Report

L2 ASCDS Version : 10.7.1

Observation 22099 - L2 Version 2
Chandra X-Ray Center

L2 Processing Date : Feb 21 2019

See axaff22099N002_VV001_vvref2.pdf for the full report

V&V Scientist	Beth Sundheim
V&V Date (YYYY-MM-DD)	2019.02.21
V&V Edition	1
V&V Disposition and Status	OK
V&V Charge Time	10.950400163174

Comments

Joint proposal with HST.

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The guide star in slot 7 was removed from the aspect solution due to poor data quality. The aspect solution is improved by the removal of this slot from the solution.

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The focal plane temperature during part of this observation was warmer than the upper limit for optimum calibration of the ACIS gain and spectral resolution (i.e., -114.0 C for ACIS-I and -112.0 C for ACIS-S).

The Chandra calibration team calibrates the ACIS gain and spectral resolution using data from the external calibration source (ECS). ECS data show that the frontside-illuminated (FI) CCDs are more temperature sensitive than the backside-illuminated (BI) CCDs.

A summary of the current calibration status of the ACIS gain and spectral resolution can be found at:

http://asc.harvard.edu/cal/Acis/Cal_prods/Gain_and_Spectral_Resolution/A_CIS_response_summary.html

The main points are:

- 1) The gain on BI chips remains within 0.3% (i.e., the systematic uncertainty in the ACIS gain quoted on the Chandra Calibration Status Summary web page) at all measured temperatures.
 - 2) The gain on FI chips remains within 0.3% below row 600 at all measured temperatures.
 - 3) The gain on FI chips above row 600 can be underestimated by as much as 1% for focal plane temperatures exceeding -116 C.
 - 4) The spectral resolution (i.e., FWHM) on BI chips is insensitive to the focal plane temperature.
 - 5) Warmer focal plane temperatures increase the FWHM on FI chips by up to 30 eV near row 512 and by up to 70 eV near the top of the chips.
- In summary, the user should be cautious in the spectral analysis of high S/N emission lines detected on the top half of FI chips in this observation. Default processing with the current version of the CALDB will underestimate photon energies by up to 1% and broaden emission lines by up to 70 eV.

seq_num	801767	Sequence number
obs_id	22099	Observation id
title	The Chandra Strong Lens Sample: Revealing Baryonic Physics In Strong Lensing Selected Clusters	Proposal title
observer	Matthew Bayliss	Principal investigator
object	SDSSJ1110+6459	Source name
dtcycle	0	
cycle	P	events from which exps? Prim/Second/Both
ra_targ	167.575	Observer's specified target RA [deg]
dec_targ	64.996389	Observer's specified target Dec [deg]
ra_nom	167.44986982609	Nominal RA [deg]
dec_nom	65.01700823039	Nominal Dec [deg]
roll_nom	155.32207891262	Nominal Roll [deg]
revision	2	Processing version of data
ontime	10950.400163174	Sum of GTIs [s]
livetime	10811.739602768	Livetime [s]
ontime0	10950.400163174	Sum of GTIs [s]
ontime1	10950.400163174	Sum of GTIs [s]
ontime2	10940.676882029	Sum of GTIs [s]
ontime3	10950.400163174	Sum of GTIs [s]
ontime6	10950.400163174	Sum of GTIs [s]
ontime7	10950.400163174	Sum of GTIs [s]
l2events	84235	Number of level 2 events

