

- **cflow: cooling flow**

A cooling flow model after Mushotzky & Szymkowiak (*Cooling Flows in Clusters and Galaxies*, ed. Fabian, 1988). An index of zero for the power-law emissivity function corresponds to emission measure weighted by the inverse of the bolometric luminosity at that temperature. The model assumes  $H_0 = 50$  and  $q_0 = 0$ . The abundance ratios are set by the **abund** command.

par1	index for power-law emissivity function
par2	low temperature (keV)
par3	high temperature (keV)
par4	abundance relative Solar
par5	redshift, $z$
norm	Mass accretion rate (solar mass/yr)